High Contrast Imaging and Wavefront Control with a PIAA Coronagraph: Laboratory System Validation

Olivier Guyon

National Astronomical Observatory of Japan, Subaru Telescope, Hilo, HI 96720

guyon@naoj.org

Eugene Pluzhnik

NASA Ames Research Center

Frantz Martinache, Julien Totems, Shinichiro Tanaka

National Astronomical Observatory of Japan, Subaru Telescope, Hilo, HI 96720

Taro Matsuo

NASA Jet Propulsion Laboratory, Pasadena, CA 91109

Celia Blain

University of Victoria, Victoria, BC, Canada V8W 3P6

Ruslan Belikov

NASA Ames Research Center

ABSTRACT

In a Phase-Induced Amplitude Apodization (PIAA) coronagraph, the telescope pupil is geometrically remapped in a gaussian-like beam by a set of two highly aspheric optics (mirrors or lenses). Thanks to this largely lossless apodization, the PIAA coronagraph offers a nearly full throughput (94.3% for the configuration tested, excluding losses due to optics coatings), preserves the angular resolution of the telescope, and can deliver high contrast very close to the optical axis. A prototype PIAA system including active wavefront control has been assembled in the laboratory at Subaru Telescope. The system operates in air with monochromatic light. Coherent light in this system has been suppressed to the dynamic speckle floor imposed by residual turbulence and vibrations, at 4.5e-8 contrast from 1.65 λ/D (inner working angle of the coronagraph configuration tested) to 4.4 λ/D (outer working angle). In addition, a static non-coherent background, likely due to multiple reflections in the system, is observed at 1.6e-7 contrast. Pointing errors are controlled at the 1e-3 λ/D level using a dedicated low order wavefront sensor. The raw contrast achieved already exceeds requirements for a ground-based Extreme Adaptive Optics system aimed at direct detection of exoplanets. In a 4 hour long exposure, we have demonstrated coherent averaging of speckles to the 3e-9 contrast level: this result is particularly encouraging for ground based Extreme-AO systems seeking long term stability to recover planets much fainter than the fast boiling speckle halo. This experiment validates technologies and algorithm which will be the core of the Subaru Coronagraphic Extreme-AO (SCExAO) system currently under assembly, and serves as a prototype for upcoming higher contrast experiments aimed at validating PIAA technologies for direct imaging of Earth-like planets from space.

Subject headings: instrumentation: adaptive optics — techniques: high angular resolution

1. Introduction

An imaging system aimed at dection or characterization (spectroscopy) of exoplanets must overcome the large contrast beween the planet and its star. This is particularly challenging for Earth-like planets, where the contrast is $\approx 10^{-10}$ in the visible and the angular separation is $0.1''$ for a system at 10pc. Many coronagraph concepts have recently been proposed to overcome this challenge (see review by Guyon et al. (2006)). Among the approaches suggested, Phase-Induced Amplitude Apodization (PIAA) coronagraphy is particularly attractive. In a PIAA coronagraph, aspheric optics (mirrors or lenses) apodize the telescope beam with no loss in throughput. A PIAA coronagraph combines high throughput, small inner working angle (2 λ/D at 10^{-10} contrast), low chromaticity (when mirrors are used), full 360 degree discovery space, and full $1\lambda/D$ angular resolution. This concept, orginally formulated by Guyon (2003), has since been studied in depth in several subsequent publications (Traub & Vanderbei 2003; Guyon et al. 2005; Vanderbei & Traub 2005; Galicher et al. 2005; Martinache et al. 2006; Vanderbei 2006; Pluzhnik et al. 2006; Guyon et al. 2006; Belikov et al. 2006; Guyon et al. 2009; Lozi et al. 2009), which the reader can refer to for detailed technical information.

In the first laboratory demonstration of the PIAA concept (Galicher et al. 2005), lossless beam apodization was demonstrated, and the field aberrations introduced by the PIAA optics were confirmed experimentally. In this first prototype, the PIAA optics unfortunately lacked surface accuracy required for high contrast imaging, and since this experiment did not include active wavefront control, the high contrast imaging potential of the technique could not be demonstrated. In the present paper, we report on results obtained with a new PIAA prototype which includes wavefront control. Our prototype combines the main elements/subsystems envisionned for a successful PIAA imaging coronagraph instrument, with the exception of corrective optics required to remove the strong off-axis aberrations introduced by the PIAA optics. This last subsystem has been designed and built for another testbed, and its laboratory performance is reported in another paper (Lozi et al. 2009).

The overall system architecture adopted for our test is presented and justified in §2. The design of the main components of the coronagraphs (PIAA mirrors, masks) is also described in this section. Wavefront control and calibration are discussed in §3. Laboratory results are presented in §4.

2. Laboratory system architecture

2.1. Coronagraph architecture

The coronagraph architecture adopted is a hybrid PIAA (Pluzhnik et al. 2006), where beam apodization is shared between the aspheric PIAA mirrors (described in §2.3) and a post-apodizer (described in §2.4). The PIAA mirrors perform most of the apodization, but leave a small amount of excess light at the edge of beam (left at 0.85% of the surface brightness at the center of the beam), which is then removed by the apodizer. Thanks to this hybrid approach, the PIAA mirrors are more easily manufacturable (less aspheric) and the apodizer tolerances are relaxed (the apodizer is not absorbing light in the bright parts of the beam). The hybrid design also solves the problem of propagation-induced chromaticity (Vanderbei 2006), which would otherwise limit contrast at $\approx 10^{-7}$ in a non-hybrid system working in a 20% wide band. The cost in throughput and angular resolution due to the apodizer are small since the apodizer only removes light in the fainter edges of the remapped beam.

A high contrast image is formed after the apodizer, where starlight is blocked by the focal plane mask. Since the upstream PIAA optics + apodizer have apodized the beam with little loss in telescope angular resolution, the focal plane mask is small, with a radius ranging from approximately 1 λ/D on the sky for a 10^{-6} contrast goal to approximately 2 λ/D on the sky for a 10^{-10} contrast goal. The focal plane mask is also part of the low order wavefront sensor (Guyon et al. 2009) briefly described in §3.2 which uses starlight reflected by the focal plane mask for accurate sensing of pointing errors and defocus.

The optical layout of the laboratory experiment is shown in Figure 1. The light source is a single mode fiber fed by a HeNe laser ($\lambda = 632.58$ nm), mounted on a x,y,z stage for control of the input tip/tilt and focus. The PIAA system (mirrors PIAA M1 and PIAA M2) creates a converg-
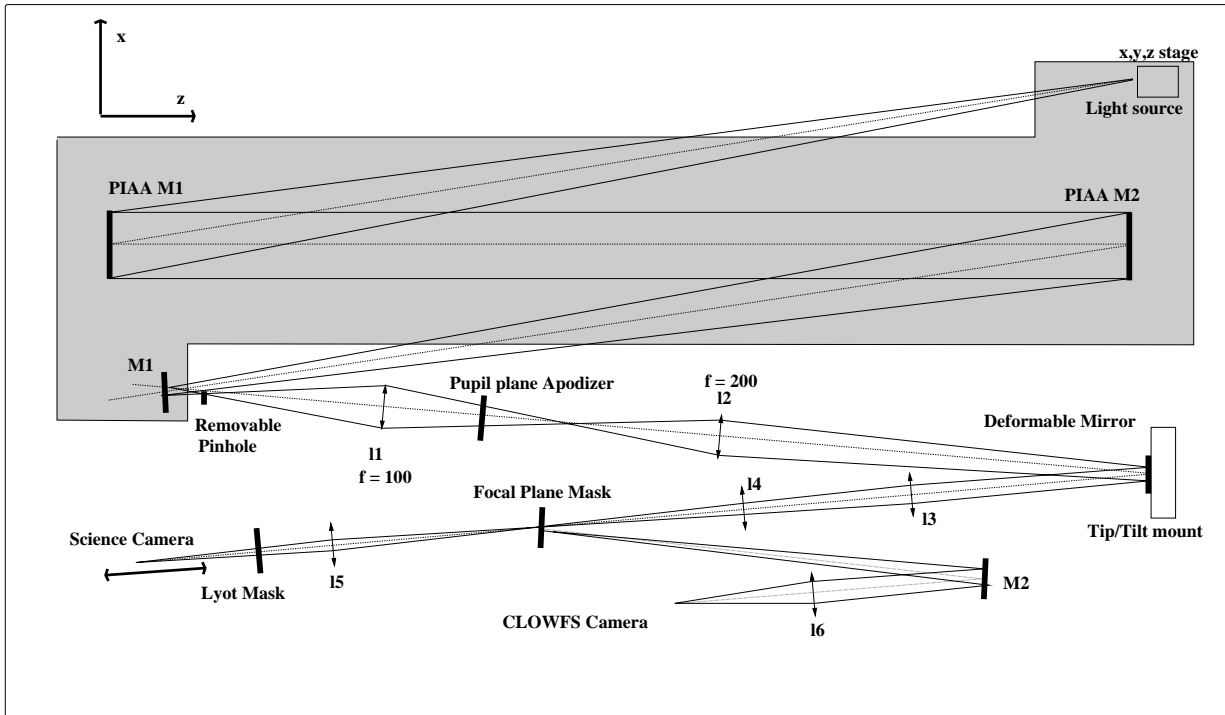
Fig. 1.— Optical layout of the laboratory PIAA coronagraph system. The grey shaded area shows the rigid PIAA bench on which the two PIAA mirrors are mounted. The light source is at the upper corner of the figure. The focal plane mask (near the bottom, center) separates light into the imaging channel and the coronagraphic low order wavefront sensor (CLOWFS) channel.

ing apodized beam. PIAA M2 is chosen as the pupil plane for the system, and lens l1 creates a small image of the pupil plane onto the apodizer. Lens l2 reimages the pupil plane on the deformable mirror, which is located in a weakly converging beam. Lenses l3 and l4 form a focal plane for the focal plane mask. The reimaging lens l5 is used to create a pupil plane and a focal plane. A Lyot mask is located in the pupil plane, but can be remotely moved out to allow the science camera (which is nominally in the focal plane) to move forward and acquire a direct pupil plane image. The focal plane mask reflects some of the light to the coronagraphic low order wavefront sensor (CLOWFS) camera. A detailed description of this device is given in Guyon et al. (2009). This design was numerically optimized to satisfy multiple constraints: size of the pupil on the deformable mirror (driven by physical size of the DM), location of critical elements on the bench driven by mechanical constraints, size of the pupil plane of science

camera detector, plate scale in the focal plane image. The same optimization code was also used to compute offsets in the position of several components during fine alignment of the system.

2.2. Deformable mirror location

The location of the deformable mirror in a PIAA system can have a large impact on the outer working angle (OWA) of the system. The OWA is defined by the furthest (from the optical axis) speckles that the DM(s) can cancel in the science focal plane. In non-PIAA coronagraph, the OWA is $(N/2) \times (\lambda/D)$, where N is the number of actuators along the diameter of the pupil. In a PIAA system, the relationship is more complicated, due to the "phase slope amplification" factor $\beta_a > 1$ introduced at the center of the beam by the remapping (see Figure 2). While in a normal imaging system of focal length f, the image of a source at angular separation (on the sky) α from the optical axis will be at position $x = f\alpha$ from the focal plane

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center, the offset will be amplified to $x = f\alpha\beta_a$ in a PIAA system. In the outer part of the beam, the phase slope has been multiplied by $\beta_r < 1$ by the remapping, explaining why in a off-axis PIAA image, some of the light stays close to the optical axis even if the main lobe of the PSF is far from the center of the image. Typical values for these quantities are $\beta_a = 3$ and $\beta_r = 0.3$.

These factors can quantiavely be used to estimate the OWA for several PIAA+DM system configurations (Figure 3):

- Configuration 1 (DM located after the PIAA optics, no inverse PIAA). The DM can control speckles up to $r = f(N/2)\lambda/D$ from the center of the science image, corresponding to an OWA equal to $\alpha = (1/\beta_a) \times (N/2) \times (\lambda/D)$ on the sky. Beyond this separation, the DM does not have sufficient actuator density at the center of the pupil: if the actuators were mapped to the entrance of the PIAA, they would be too large at the center of the beam.

- Configuration 2 (DM located before the PIAA optics, no inverse PIAA). If the actuators were mapped to the output of the PIAA, we find that the actuator size at the edge of the apodized beam is divided by $\beta_r < 1$ (the actuators are bigger), while they are considerably smaller in the center of the apodized beam. In the outer part of the beam, the DM sampling can only create a dark hole of radius $r = \beta_r \times (N/2) \times (\lambda/D)$, corresponding to $OWA = (\beta_r/\beta_a) \times (N/2) \times (\lambda/D)$ on the sky.

In both configurations 1 and 2, the unaberrated field of view is limited to $r \approx 5\lambda/D$ by the absence of inverse PIAA optics. There is therefore little advantage to increasing the OWA much beyond this radius. We note that the OWA and the unaberrated field of view are equal for $N \approx 30$ and $N \approx 100$ in configurations 1 and 2 respectively. We now explore configurations including inverse PIAA optics to provide a wide unaberrated field of view. In these configurations, the field of view for high contrast observations is limited by the OWA of the wavefront control system.

- Configuration 3 (DM ahead of PIAA optics). In this configuration, the inverse PIAA can-

cels the effect of the forward PIAA, and the OWA is equal to what it would be in a non-remapped system: $OWA = (N/2) \times (\lambda/D)$. The only role of the remapping is to create an intermediate step where starlight is efficiently removed by the focal plane mask.

- Configuration 4 (DM after PIAA optics). The OWA can be found by remapping the DM geometry in the plane ahead of the forward PIAA, where its actuators will be magnified by β_a at the center of the pupil. The OWA, defined by the largest actuator in this plane, is therefore $OWA = (1/\beta_a) \times (N/2) \times (\lambda/D)$.

In order to optimize the use of a given number of actuators, the DM(s) should therefore be placed after the PIAA optics in a PIAA system without inverse PIAA optics (configuration 1), or before the PIAA optics if inverse PIAA optics are included (configuration 3). The configuration adopted in our laboratory demonstration is configuration 1 (DM after PIAA optics, no inverse PIAA optics).

2.3. Aspheric PIAA mirror design and fabrication

The geometric remapping is performed by two highly aspheric mirrors, the PIAA mirrors. The role of the first mirror is mostly to project on the second mirror the desired amplitude profile, which is gaussian-like with a faint "plateau" on the outside of the beam (see Figure 4, left). This PIAA M1 mirror acts as a converging element in the center (to concentrate more light in the center of the beam on PIAA M2) while light in the outside is diluted in a wide area of the beam on PIAA M2. This behavior explains the peculiar aspheric sag shown in Figure 4 on the right. The PIAA M2 mirror's role is to re-collimate light to output a beam which is apodized but free of phase aberrations.

PIAA mirrors can be designed by solving a relatively simple differential equation when the input and output beams are collimated (the equation is given in Guyon (2003), and also in Traub & Vanderbei (2003) in a different form) or when the system is on-axis. In our laboratory experiment, the PIAA mirrors are focusing elements and the aspheric remapping shapes are added to off-axis
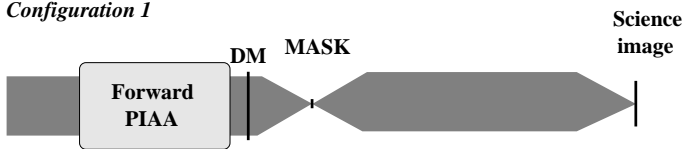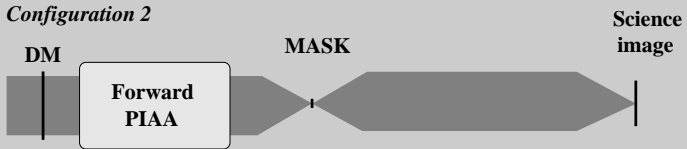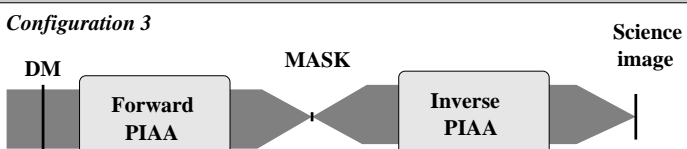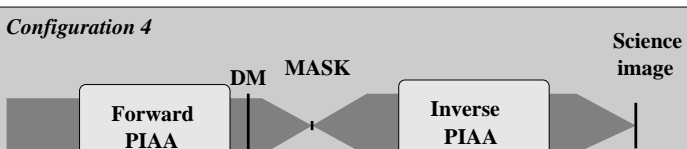
	Wavefront Control Outer Working Angle (OWA)	Unaberrated Field of View
Configuration 1 — Forward PIAA, DM, MASK, Science image	OWA = (1/3) x (N/2) x (λ/D)	~ 5 λ/D
Configuration 2 — DM, Forward PIAA, MASK, Science image	OWA = 0.1 x (N/2) x (λ/D)	~ 5 λ/D
Configuration 3 — DM, Forward PIAA, MASK, Inverse PIAA, Science image	OWA = 1.0 x (N/2) x (λ/D)	no field aberration
Configuration 4 — Forward PIAA, DM, MASK, Inverse PIAA, Science image	OWA = (1/3) x (N/2) x (λ/D)	no field aberration

Fig. 3.— Four possible architectures for a PIAA coronagraph with wavefront control. For each configuration, the outer working angle of the wavefront control system and the field of view imposed by remapping are given. The PIAA slope amplification factor $\beta_a = 3$ and slope reduction factor $\beta_r = 0.3$ are considered here. Configurations shown in gray (configurations 2, 4 and 5) should be avoided (see text for details).
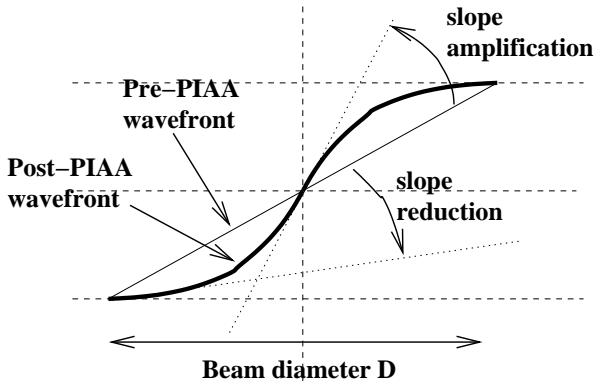

Fig. 2.— Slope amplification and reduction factors in a PIAA system. The remapping introduced by a PIAA system amplifies the wavefront slope at the center of the apodized beam (where most of the light is located) and reduces the wavefront slope at the edges of the beam.

parabolas. In this configuration, the PIAA mirror shapes cannot be derived from a simple differential equation, and they were designed by an iterative algorithm:

1. Initialization: The PIAA mirror shapes are computed by solving the differential equation for an on-axis system.

2. A constant slope is added to each of the PIAA mirror. If there were no apodization, the mirrors obtained in step 1 would be on-axis parabolas, and this slope would turn them into off-axis parabolas. In the PIAA system, however, adding this slope only leads to an approximation to the off-axis PIAA system

3. A 3 dimensional raytracing code is used to compute the beam phase and amplitude on the surface of a sphere centered on the out-
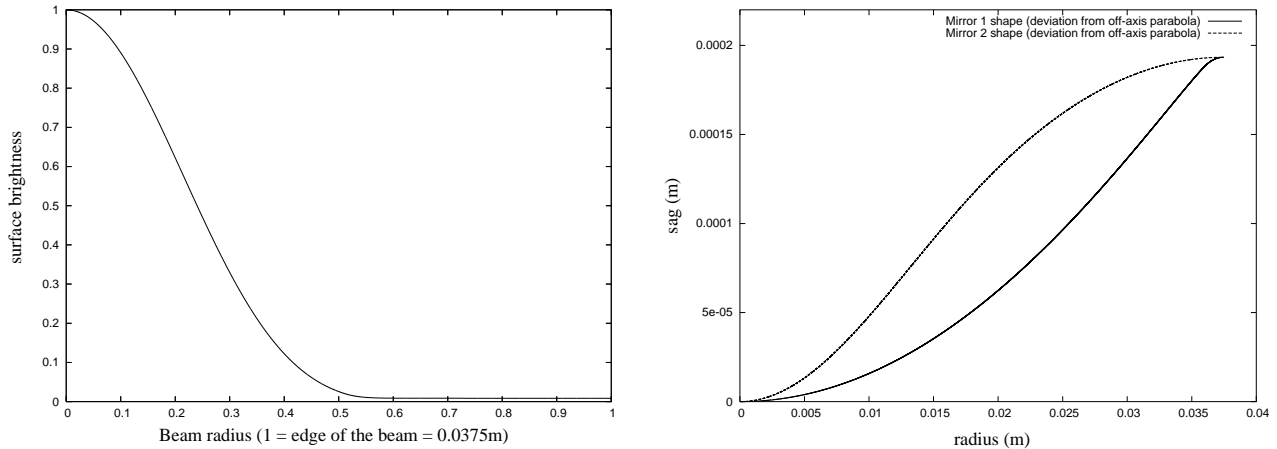
Fig. 4.— Left: Beam apodization profile used for the design of the PIAA system. In this hybrid design, the apodization is not complete, and the beam surface brightness at the edge of the beam is left at 0.85% of the center surface brightness. This extra light will need to be removed by a conventional apodizer. Right: Apodization radial sag term for PIAA M1 and PIAA M2 in an on-axis configuration. The narrow region at the edge of PIAA M1 has a strong localized curvature, and is the most challenging feature for manufacturing the PIAA system optics.

put focus of the system immediately after reflection on PIAA M2.

4. The difference between the measured and desired beam amplitude on the sphere is used to update PIAA M1's shape by linear decomposition of this residual (using pre-computed residuals obtained by adding Zernike polynomials on PIAA M1's shape).

5. The residual phase error measured on the sphere is compensated by changing PIAA M2's shape.

6. Return to Step 3 with the new mirror shapes.

This algorithm converges because changing PIAA M2's shape has little effect on the amplitude profile of the beam on the sphere, which is almost entirely a function of PIAA M1's shape.

The PIAA shapes were computed for a 75mm beam diameter at the PIAA mirrors, a 1.125m separation from the center of PIAA M1 to the center of PIAA M2, and a 190mm offset between the PIAA M1 to PIAA M2 centerline and the input and output of the PIAA system (see Figure 1). In the coordinate system shown in Figure 1, each mirror shape can be written as:

$$z(x,y) = OAP(x,y) + sag(r) + \Sigma_i^j \alpha_i Z_i^j \qquad (1)$$

- $OAP(x,y)$ is the off-axis parabola which would be the PIAA mirror shapes if no apodization was performed. It is a 1133 mm focal length OAP with a 190 mm off-axis distance from the center of the optical element. This shape is identical for the two PIAA mirrors although the orientation is different.

- $sag(r)$ is the apodization radial sag on each mirror. It is computed for an on-axis system.

- A corrective term is added to account for the fact that the system is off-axis (tends to 0 for an on-axis system), decomposed as a sum of Zernikes polynomials up to radial order 7. For both PIAA mirrors, this correction is ≈95 nm RMS (excluding tip-tilt).

The most challenging feature of the system is the small radius of curvature in the outer part of the beam on PIAA M1. While the hybrid design adopted mitigates this problem, the radius of curvature still reaches a minimum of 155 mm near the edge of the mirror, at 36.1mm from the center of the mirror.

The PIAA mirrors were fabricated by Axsys Imaging Technologies. The mirror substrates were initially diamond turned according to the 3-D prescriptions described above, and then polished

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against computer generated holograms (CGHs). PIAA M1 and PIAA M2 were then assembled on a rigid aluminum bench, aligned and permanently fixed to the bench. The residual system wavefront error was then reduced to 0.04 waves RMS by figuring PIAA M2. Two sets of PIAA mirrors (4 mirrors total) were manufactured.

2.4. Post-apodizer and system throughput

In order to ease manufacturing, the PIAA optics were designed to perform most, but not all, of the beam apodization required for high contrast imaging. A more conventional apodizing scheme is therefore necessary to transform the beam profile at the output of the PIAA optics (solid curve in Figure 5) into the desired beam profile (dashed curve in Figure 5).

The post apodizer was designed in transmission, with a series of narrow opaque rings blocking light. The position and width of the rings is optimized to best approximate the ideal continuous apodization profile shown in Figure 5 as the "Apodizer transmission" curve. Several contraints were imposed on the design to ensure manufacturability: no ring should be less than 0.8 μm wide and the gap between consecutive opaque rings should be no less than 5 μm. The resulting design is composed of 109 opaque rings for a total apodizer diameter of 3.815 mm (defined by the outer edge of the last opening between opaque rings). The apodizer was manufactured by lithography on a transmissive substrate.

The post-apodizer throughput over the 3.815mm diameter is 96.9 %, but due to the narrow rings in the apodizer, some of the light transmitted is diffracted at large angles. The "effective" throughput of the apodizer is 94.3%, and would be equal to the "raw" throughput if the apodizer were continuous instead of binary. The full system throughput can therefore reach 94.3% (excluding losses due to coating) provided that the telescope pupil size on PIAA M1 is adjusted to the apodizer diameter. In practice, the telescope pupil should however be made slightly larger to allow for pupil centration errors, and in very high contrast applications (space coronagraphy), to mitigate possible edge "ringing" effects due to Fresnel propagation.

The apodizer throughput was mesured by inserting a pinhole in the PIAA output focus and
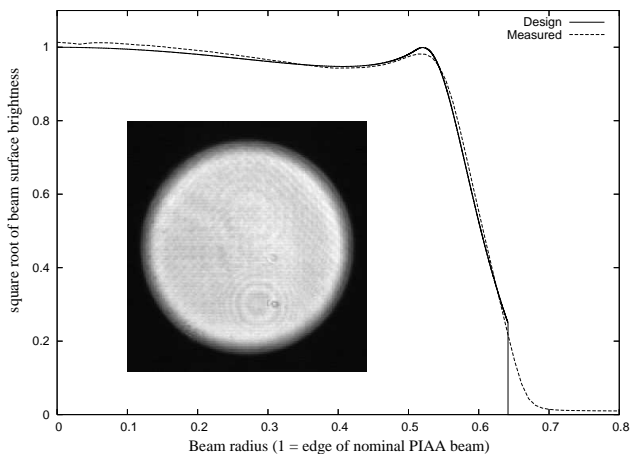

Fig. 6.— Laboratory measurement of the apodizer throughput. The designed and measured radial throughputs are compared.

moving the science camera in the pupil plane. The 1 μm pinhole is used to de-apodize the beam at the expense of a very low throughput. As shown in Figure 6, the measured apodizer profile agrees very well with the designed apodization. The residual difference between the two curves is due to the finite size of the pinhole (the beam before the apodizer is slightly apodized, so the measured profile is slightly too bright in the center) and the finite angular resolution of the pupil re-imaging (the sharp edge of the apodizer is blurred).
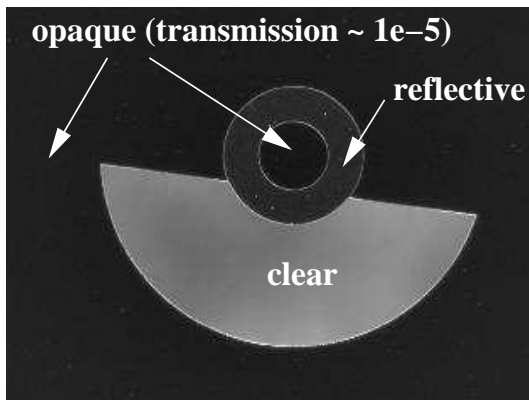
2.5. Focal plane and Lyot masks

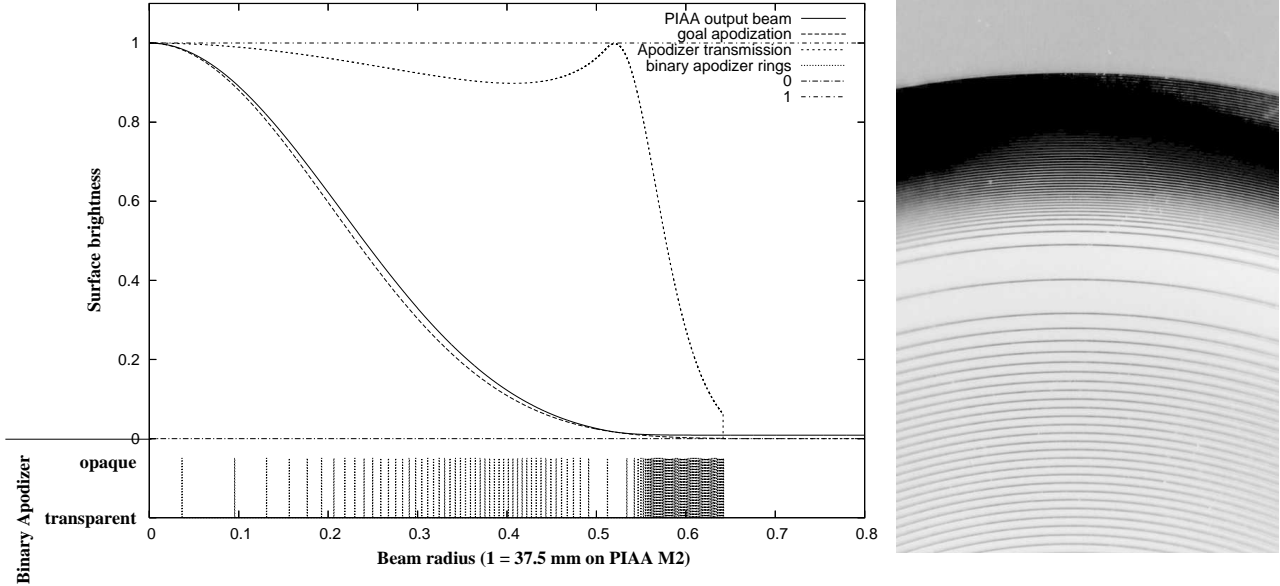Fig. 7.— Microscope image of the focal plane mask

Fig. 5.— Left: Apodizer design. Right: Microscope image of the outer part of the apodizer (the image covers approximately 1mm vertically). The outer edge of the apodizer (last transmissive ring) is at 1.9mm radius. The width of the individual opaque rings is $0.8\mu m$.

The focal plane mask, shown in Figure 7, is used in transmission, and serves three purposes:

- Blocks the bright central PSF core. The radius of the non-transmissive central zone of the focal plane mask defines the inner working angle of the coronagraph, which is 1.65 λ/D in our experiment.

- Transmits the science field to the science camera. The shape of the clear zone of the focal plane mask is chosen to exclude regions of the focal plane where the wavefront control system cannot remove diffracted light. Since our experiment uses a single deformable mirror, diffracted light can only be controlled over half of the field of view. The clear opening in the mask is therefore "D"-shaped with an outer radius imposed by the DM actuator sampling. A slightly larger rectangular zone could also have been adopted, but would have imposed the rotation angle of the focal plane mask.

- Reflects some of the starlight to the low order wavefront sensor (LOWFS) camera. A reflective annulus, extending from 0.8 λ/D

to $1.65\lambda/D$, is used for this purpose.

The "opaque" zones of the mask are not fully opaque: their transmition and reflection are respectively $\approx 10^{-5}$ and $\approx 10\%$.

A Lyot mask is located in the pupil plane between the focal plane mask and the science focal plane. This mask is designed to block all light outside the geometrical pupil and transmit all light within the pupil. It therefore has no effect on the nominal system throughput, and its role is to ensure that the scattered light reaching the focal plane camera does not contain light outside the pupil. Although correcting for such light is theoretically possible if it is coherent, it requires an accurate model of the coronagraph which can predict how light outside the pupil is affected by DM actuator positions. The Lyot mask was made by drilling a small hole in an aluminum plate, and its diameter is slightly smaller than the pupil size to account for alignment tolerances.

Both the focal and Lyot masks are on motorized stages and can be removed from the beam.

3. Wavefront Control

3.1. Initial Calibration Loop (without focal plane mask)

Initial calibration is performed using conventional phase diversity with no focal plane mask: images are acquired with the science camera in six positions regularly spaced from the focal plane to the pupil plane. An iterative Gerchberg-Saxon algorithm is used to reconstruct the pupil plane complex amplitude. As shown in Figure 8 (top left), the beam quality is intially quite poor, with a large amount of astigmatism. The corresponding focal plane image is shown in Figure 8, bottom left.
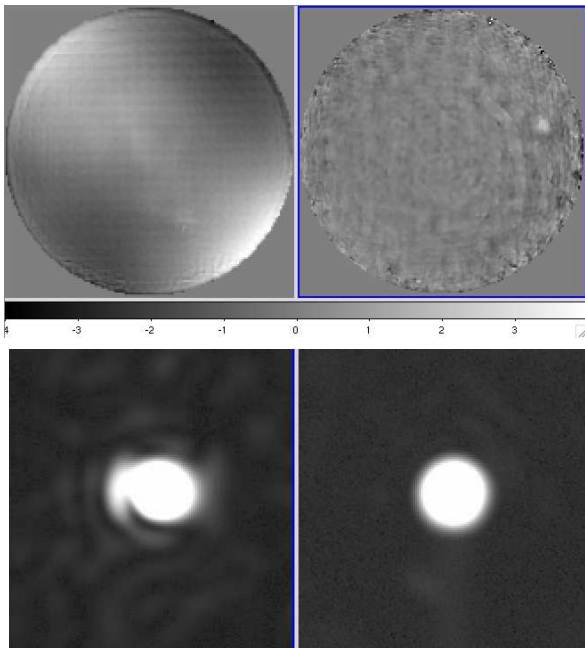

Fig. 8.— Pupil phase (top) and focal plane image (bottom) when the deformable mirror is powered off (left) and set to its nominal position after calibration (right). A malfunctionning actuator is visible on the right side of the beam.

The phase diversity routine described above is reapeated N times ($N \approx 10$), with a different set of DM voltages applied for each phase diversity measurement sequence. The N phase maps obtained and the N DM voltage maps used to obtain them are then used to constrain a model of the DM response which consists of seven parameters:

- Geometrical correspondance between the DM and the pupil image: x and y shift, scale, and rotation

- Physical constants describing the DM behavior: width of the actuator influence function, DM dispacement for 100V applied and power index α in the displacement to voltage relationship (displacement $\propto V^{\alpha}$).

The result of the DM calibration can then be used to flatten the wavefront measured and produce a sharp focal plane image (Figure 8, right).

This initial calibration is a necessary preliminary step for the high contrast wavefront control, which needs (1) a knowledge of the starting point (typically less than 1 radian error on the wavefront) and (2) a good understanding of how DM commands affect the pupil plane phase.

3.2. Low order wavefront errors

Low order wavefront errors are measured by reflecting a portion of the bright starlight masked by the coronagraph focal plane into a dedicted camera. A detailed description of this low order wavefront sensor (LOWFS) can be found in Guyon et al. (2009). The LOWFS signal is used to simultaneously drive the deformable tip-tilt and the source position ahead of the PIAA optics. A key feature of the LOWFS is the ability to separate pointing errors (pre-PIAA tip-tilt) from post-PIAA tip-tilt, which is essential to maintain high contrast: even a small pre-PIAA tip-tilt creates diffraction rings outside the IWA of the coronagraph, and pre-PIAA tip-tilt errors cannot be compensated for by post-PIAA tip-tilt.

When the low-order loop is closed, the measured residual pointing error is $10^{-3}\lambda/D$, and is therefore small enough to be negligible in the scattered light error budget shown in §4.3. A more detailed description of the design, calibration, control algorithm and performance of the LOWFS in our experiment is given in Guyon et al. (2009).

3.3. High order wavefront control loop

Coherent scattered light in the "clear" opening of the focal plane mask is measured by phase diversity introduced on the DM. A series of focal plane images, each acquired with a slightly different DM
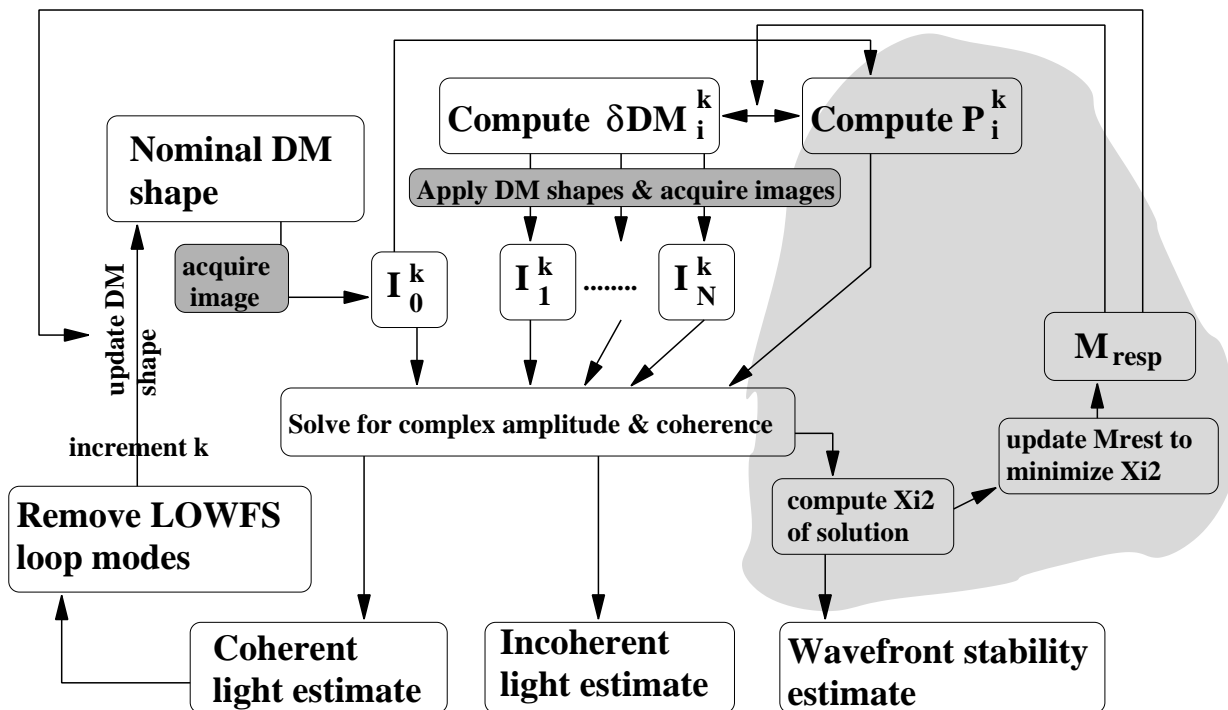
Fig. 9.— High order wavefront control loop, showing both the main loop and the system response matrix optimization loop (light shaded area). The two dark shaded boxes indicate image acquisition, which in the "simulation" mode, can be replaced with a simulated image acquisition using a model of the experiment and DM response.

shape, is used to reconstruct the complex amplitude and coherence of the scattered light. The high order wavefront control loop uses a linearized representation of the system in focal plane complex amplitude, as described in the electric field conjugation (EFC) approach proposed by Give'on et al. (2007). The wavefront control loop is shown in Figure 9, and is built around the EFC approach.

Prior to starting the loop, a model of the coronagraph is used to compute how each actuator motion affects the complex amplitude in the focal plane. Since this relationship is linear for small displacements, this model is stored as a complex amplitude system response matrix M_{resp} (shown on the left part of Figure 9) of size n by m, where n is the number of DM actuators (ignoring actuators outside the pupil) and m is the number of pixels in the high contrast region of the focal plane.

3.3.1. Loop initialization

For each iteration k of the loop, the first step in the wavefront sensing process is to acquire an image I_0^k with the DM shape set at the best known position for high contrast imaging (upper left corner of Figure 9). This first image is then used to choose the shapes to apply on the DM to optimize measurement accuracy and sensitivity.

3.3.2. Compute wavefront sensing DM shapes

For the control loop iteration k, we denote $\delta DM_i^k(u,v)$, with $i = 1...N$, the N DM displacement (compared to the "reference" position of the DM), and $P_i^k(x,y) = M_{resp}\delta DM_i^k(u,v)$ the corresponding complex amplitude "probes" which are added to the focal plane by each of these DM offets. As a guideline, it is best to choose these DM offsets so that the additional light (the complex amplitude focal plane "probes") is approximately as bright as the light which needs to be measured. If the probes are too strong, the measurement is

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too sensitive to errors in the DM calibration; if they are too weak, the measurement is contaminated by photon noise, readout noise and small variations in the incoherent scattered light. Finally, randomly modulating the probes can mitigate the effect of calibration errors.

- The first probe P_1^k is chosen to satisfy, for each pixel (x, y):

$$|P_1^k(x, y)|^2 = \alpha_0 + \alpha_1 I_0^k(x, y) \qquad (2)$$

where $I_0^k(x, y)$ is the image acquired with DM shape DM_0^k. If $\alpha_1 = 0$, this constraint will force the DM shape to add a uniform coherent background of contrast α_0 in the focal plane, while if $\alpha_0 = 0$ and $\alpha_1 = 1$ it will drive the DM shape to add a speckle map with the same intensity as in the I_0^k image. The phase of this probe is not constrained, and is chosen by an iterative to best satisfy equation 2 using M_{resp}, with δDM_1^k as the free parameter. Since there is ususally no exact solution to this equation, the DM command which best satifies this criteria are chosen with a regularisation parameter to prevent excessive DM displacements.

- The second probe is chosen so that, at each point (x, y) in the focal plane, its amplitude is identical to the first probe, but its phase is offset by $\pi/2$:

$$P_2^k(x, y) \approx i\, P_1^k(x, y) \qquad (3)$$

This $\pi/2$ phase offset maximizes the WFS sensitivity if the dominant sources of noises are photon noise and readout noise (Guyon et al. 2005). We note that if all DM actuators are functionning, there is a perfect solution to this equation, which can be obtained by shifting each spatial frequency of the DM_1^k map by $\pi/2$. Images acquired with these first two probes, together with the image I_0^k, would be sufficient to solve for wavefront errors if light in the focal plane if fully coherent, but at least one more probe is needed to unambiguously measure light coherence, and more probes can also provide the redundancy required for implementation of the diagnostic tools described in §3.3.3 and §3.3.4.

- Two additional probes can be chosen to be $P_3^k = -P_1^k$ and $P_4^k = -P_2^k$, with exact solutions $\delta DM_3^k = -\delta DM_1^k$ and $\delta DM_4^k = -\delta DM_2^k$ respectively.

- In our laboratory experiment, we chose to also add 5 more probes with random uncorrelated DM shapes of similar amplitude than the DM displacements obtained for probes 1 to 4 above.

3.3.3. Solving for complex amplitude and coherence

For each pixel (x, y) of the focal plane detector, the complex amplitude $A(x, y)$ of the coherent light leak and the intensity $I(x, y)$ of the incoherent light leak is estimated by solving the following set of equations:

$$I_i^k(x, y) = |A(x, y) + P_i^k(x, y)|^2 + I(x, y) \qquad (4)$$

for $i = 0...N$, with $P_0^k(x, y) = 0$ (image acquired with nominal DM shape).

This set equation has three unknowns and is therefore overconstrained for $N > 2$. With $N = 9$ adopted in our experiment, we can measure the fit quality by computing the fit χ^2, which includes errors due to:

- Photon and readout noise in the $I_i^k(x, y)$ measurements

- Variations in the light leaks during the measurement sequence. When solving for this set of equation, we assume $A(x, y)$ and $I(x, y)$ are static, but if they vary, χ^2 will increase.

- Errors in M_{resp}, which lead to errors in the estimation of the values of $P_i^k(x, y)$. If M_{resp} is wrong, then the DM command sent will not produce the expected $P_i^k(x, y)$.

For convenience, we have scaled χ^2 in coronagraphic contrast unit. This scaling is performed by measuring the uncorrelated "noise" that would need to be added to $A(x, y)$ between frames to reproduce the observed value of χ^2. In this unit, the observed χ^2 is approximately $4.5\,10^{-8}$ (see §4).

The first contribution (photon + readout noise) has been computed to be a small part of the χ^2 observed. We observed that increasing α_1 (see equation 2) above 1 has little effect on the residual χ^2,

also independantly suggesting that χ^2 is not due to detector readout noise.

3.3.4. System response matrix optimization

One key output of the wavefront control loop is the estimation of coherent light leaks (which should be used to compute the DM correction to apply for the next iteration), incoherent light leaks (which the DM can do nothing about) and the measurement of the wavefront stability during the measurement sequence. All these quantities depend upon a reliable estimation of M_{resp}. It is for example possible, if M_{resp} is wrong, to obtain a low value of the residual coherent light and think the system has converged to a good contrast value, while in fact a significant amount of coherent light remains. This last issue is mitigated, but not entirely addressed, by continuously varying the probes (as this error is a function of the probes chosen, and will average to zero if the probes are "randomly" chosen). An error in M_{resp} would first appear as a large χ^2 value for the solution of equations 4.

To address this, we have added a M_{resp} optimization loop within our control loop. For each iteration k, the derivative of χ^2 with M_{resp} is computed (this is a total of $2 \times n \times m$ derivatives, as the derivative is computed for the real and imaginary parts of each element of the M_{resp} matrix). With careful regularization, M_{resp} is then slightly modified in order to reduce the sum of the χ^2 over all pixels of the image. This algorithm was first tested on simulated data with an initial M_{resp} estimate which was different from the actual M_{resp} used in the simulation for computing the images. This test showed that M_{resp} did converge toward the "true" M_{resp}, and that the χ^2 value decreases as a result. Convergence is very slow, due to the large number of coefficients in the M_{resp} matrix, requiring several hundred iterations before a significant improvement in χ^2 is observed.

3.3.5. Correction applied to the DM

The linear electric field conjugation (EFC) algorithm (Give'on et al. 2007) is used to cancel coherent scattered light. This algorithm uses the linearized coronagraph model which is also used for the measurement step described above. The system response matrix is inverted to build a control matrix which is multiplied to the coherent light estimate to yield the DM shape offset to be applied. Regularization schemes proposed by Give'on et al. (2007) were used to improve the loop stability and convergence speed.

4. Laboratory results

4.1. Alignment and Apodization measurement in the pupil plane

The pupil plane apodization map is measured by placing the science camera in the pupil plane. In our experiment, the pupil is conjugated simultaneously to the PIAA M2 mirror, the apodizer, and the DM. Alignment is necessary to ensure that these three planes are conjugated and that their relative scales are correct. The camera positions for which conjugation to these planes is achieved are measured and the corresponding pupil scales are derived from the images. These six numbers are then fed to an optimization routine which computes the offsets to be applied to all movable optical elements after PIAA M2 to meet the conjugation and scale requirements. A few iterations of this sequence were sufficient to converge.

In the fine alignment step, the apodizer alone is moved. The pupil image is compared to a simulated pupil image where the relative scale and lateral offset between the PIAA apodization and the apodizer transmission map are free parameters. The values of these three parameters which give the smallest residual difference is then used to guide fine alignment of the apodizer. Fine tuning of the scale between the apodizer and the PIAA apodization is possible because the beam at the apodizer is non collimated: apodizer motion along the optical axis changes this scale. As shown in Figure 10, the pupil apodization profile measured after alignment is in good agreement with the theoretical profile.

4.2. Imaging with a non-corrected PIAA system

Figure 11 shows the system on-axis PSF in "imaging" mode (no coronagraph focal plane mask). The on-axis PSF is similar to an Airy function without the Airy rings beyond 1.22 λ/D.

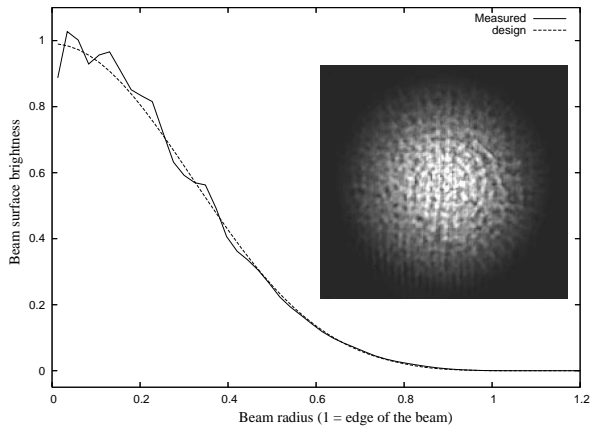While the on-axis image is sharp and exibits high contrast, our laboratory system did not in-

Fig. 10.— Measured apodization profile, compared the to the beam profile the experiment was designed to deliver. A 2-D image of the beam is also shown.

clude the "inverse" PIAA optics necessary to correct for the strong off-axis aberrations introduced by the "forward" PIAA optics. These inverse optics do not need to be of coronagraphic quality, and can be a small set of lenses. A laboratory demonstration of wide field correction with inverse PIAA optics is described in a separate paper (Lozi et al. 2009).

Without PIAA optics, the image of an off-axis source rapidly changes shape as the source moves away from the optical axis. As shown in Figure 12, measured off-axis PSFs are in good agreement with simulations using a remapping of the beam phase (Guyon 2003). With such a strong field aberration, measuring the focal plane plate scale is challenging and its value is a function of the metric used. In this paper, we choose to adopt the non coronagraphic PSF photocenter to measure plate scale. In Figure 11, the $4\,\lambda/D$ ring therefore shows where the photocenter of the PSF would be if the source was $4\,\lambda/D$ from the optical axis of the entrance telescope. The bright PSF core at this separation would be slightly outside the ring, but the fainter assymetric diffraction arcs of the off-axis PSF would be inside the ring. The same photocenter metric is used to measure the angular sizes on the focal plane mask given in §2.5 .
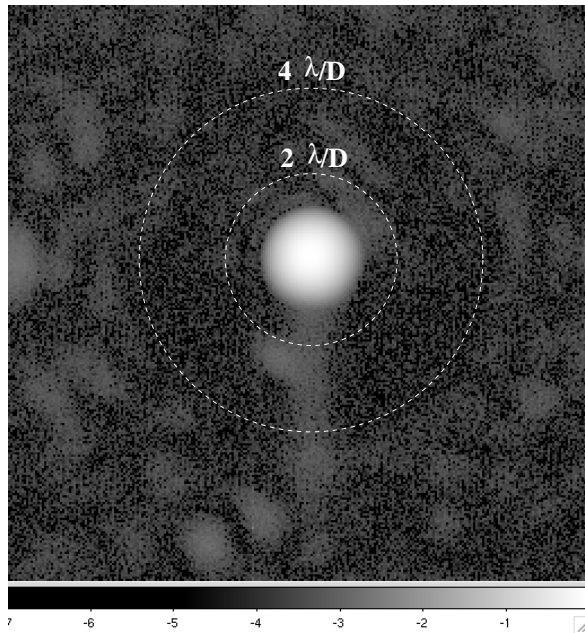

Fig. 11.— On-axis PSF without focal plane mask (log scale indicated at the bottom of the figure). The vertical "bleeding" feature extending downward of the PSF core is a camera artefact, and is removed when the focal plane mask is used. A faint ghost due to the entrance window of the camera is visible just beyond $2\,\lambda/D$.

4.3. Coronagraphic performance: measurements

Coronagraphic performance is measured with the focal plane and Lyot masks in the beam and both the LOWFS loop and high order wavefront control loop closed. Figure 13 (upper left) shows a raw image from the science camera. Most of the light reaching the camera is due to partial transmittance (at the 10^{-5} level) of the core of the focal plane mask, which produces a central peak in the image. The clear $1.65\lambda/D$ to $4.4\lambda/D$ opening in the focal plane mask is visible below this central peak.

Residual light is decomposed in two components by the wavefront sensing sytem, using the approach described in §3.3:

- An incoherent component composed of light which does not interfere with light extracted directly from the central PSF core. This component appears to be mostly stable in
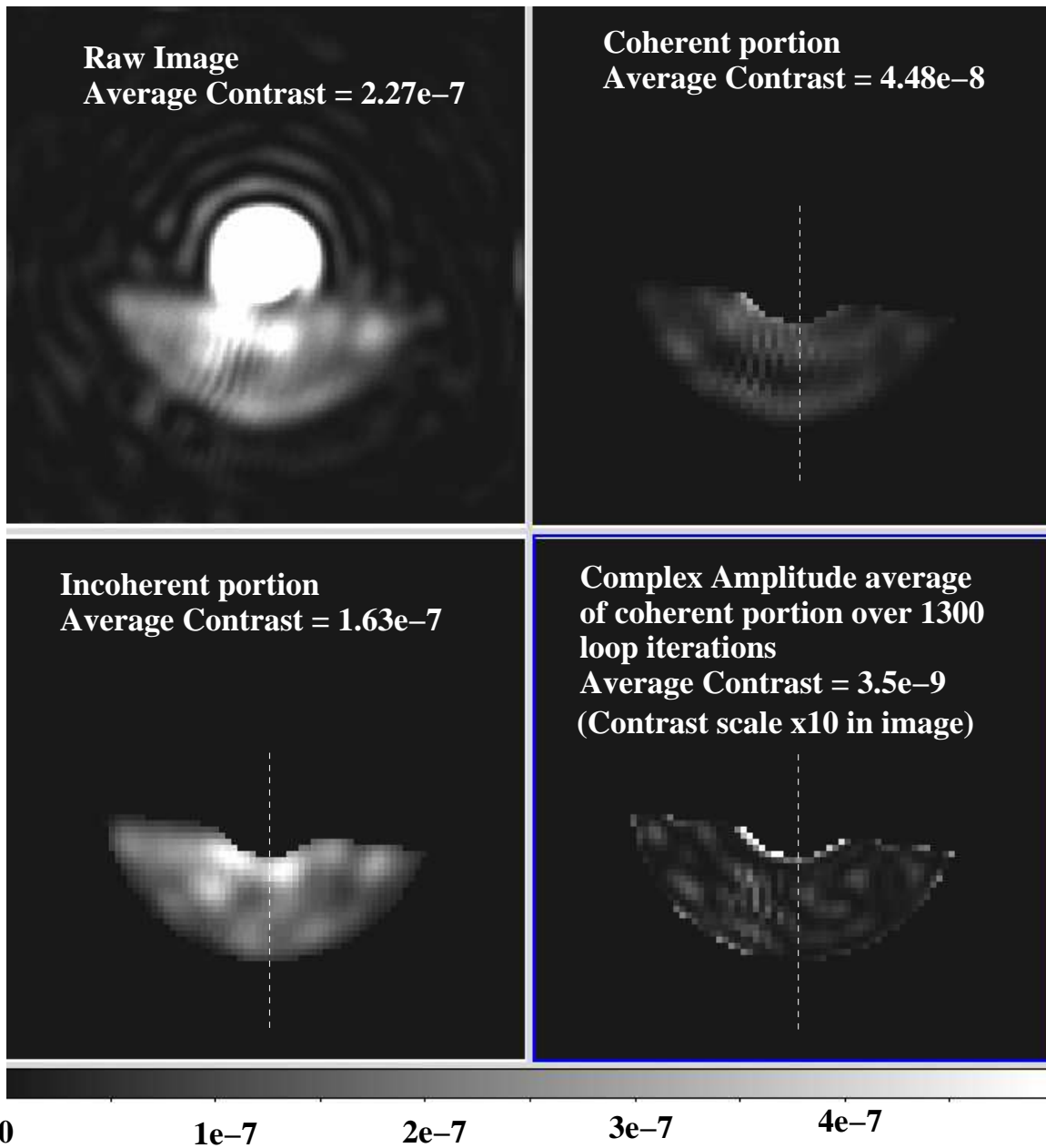
13

Fig. 13.— Top left: Raw coronagraphic image. A decomposition of the scattered light into a coherent component (top right) and incoherent component (bottom left) shows that the raw contrast is dominated by incoherent light. The coherent "bias" over a long period of time is shown in the bottom right.
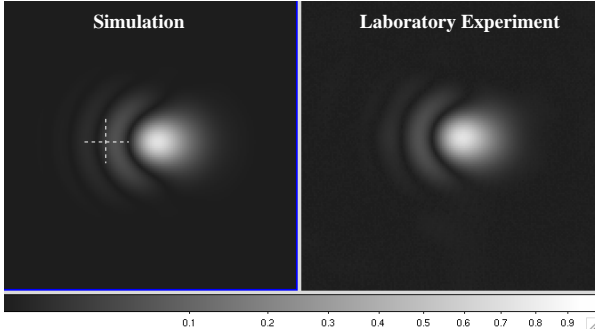
14

Fig. 12.— Off-axis PSF without focal plane mask, for a 1.7 λ/D off-axis distance. A cross indicates the optical axis in the simulated image (left). The image obtained in the laboratory (right) is in very good agreement with the simulation. Both images are shown to the same brightness scale (bottom).

structure and varies from iteration to iteration between $1.5\,10^{-7}$ and $2\,10^{-7}$ in contrast.

- A coherent component which is used to compute the DM shape for the next iteration. This component is at approximately $5\,10^{-8}$ contrast is well decorrelated on timescales above the response time of the wavefront control loop. The estimation of this component varies greatly from iteration to iteration, and is sometimes below $1\,10^{-8}$ contrast.

We also measured the χ^2 of the solution, and found it to correspond to a change in coherent light at the $4.5\,10^{-8}$ contrast level during the measurement sequence, which is 10s long. The M_{resp} optimization loop described in §3.3.4 gave no noticable improvement in χ^2, suggesting that the inital calibration led to a good estimate of M_{resp} and that the observed χ^2 is indeed dominated by fluctuations in coherent light at the $4.5\,10^{-8}$ contrast level.

Visual inspection of the coherent and incoherent portions of the light strongly suggests that this decomposition was successful:

- The images obtained show very little high spatial frequency noise. Although the reconstruction is performed independantly for each pixel, both the coherent and incoherent components show speckles and features covering several pixels in these oversampled images.

- The ghost to the lower left of the optical axis was properly analysed by our algorithm. This ghost, also visible in figure 11, is due to internal reflection in the window of the CCD camera. Because the two surfaces of the window are not perfectly parallel, this ghost is "fringed". Although the fringes are coherent, they are too thighly spaced for the wavefront control system to remove: they are the only known feature in the image that should be seen in the coherent image, should not be seen in the incoherent image, and cannot be removed by the wavefront control system. As shown in figure 13, our coherent/incoherent analysis correctly identifies these fringe as coherent, and only a minute fraction of the fringes is present in the incoherent portion of the image (likely due to small motion of the fringes during the measurement sequence).

4.4. Analysis

As summarized in table 1, together, these measurements show that:

- During the measurement sequence, coherent light is varying by $\approx 4.5\,10^{-8}$ in contrast due to turbulence or vibrations in the system, as shown by the χ^2 analysis.

- The raw contrast is dominated by a very stable incoherent component which is most likely a ghost. Given the high number of air-glass surfaces (twenty), including some which are not anti-reflection coated, this result is not too surprising. The variations observed in the estimate of the incoherent light are due to the coherent light variation during the measurement, which affects the incoherent estimate. The data obtained is compatible with a fully static incoherent background, as would be expected from a ghost.

- The coherent light leak estimate is $\approx 5\,10^{-8}$, which is at the level expected from the $4.5\,10^{-8}$ variations shown by the χ^2 analysis. The large variation, from iteration to iteration, observed in the coherent light residual is due to the turbulence/vibrations in the system. We note that the "lucky" iterations

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TABLE 1
CONTRAST BUDGET (AVERAGED OVER 1.6 TO 4.4 λ/D)

Term	Origin	Value	Calibration
Incoherent light, static	optical ghost	$1.63\ 10^{-7}$	Incoherent portion of WFS data
Coherent, variable in $t < 10s$	turbulence, vibrations	$4\ 10^{-8}$	residual from WFS reconstruction
Coherent residual (in 10s)		$5\ 10^{-8}$ (typical)	Coherent portion of WFS data
		$7\ 10^{-9}$ (best)	
Coherent, static (in 4hr)	uncorrectable coherent light	$< 3\ 10^{-9}$	time-averaged coherent light

where the coherent light is estimated below 10^{-8} are artefacts of the time averaging during the measurement period: even during these "lucky" periods the coherent light did vary by $\approx 4.5\ 10^{-8}$.

- The wavefront control loop successfully removes static coherent speckles. Over a 4hr period of time, we have measured the static coherent speckles to be at the $3\,10^{-9}$ contrast level. Except for a known ghost on the camera window, we could not detect any residual bias in the residual coherent light.

5. Conclusion

The results obtained in this experiment are especially encouraging for ground-based coronagraphy. The $2\ 10^{-7}$ raw contrast we have achieved already exceeds by two orders of magnitudes the raw contrast that can be hoped for in even an theoretically ideal Extreme-AO system (Guyon et al. 2005). We note that with a more careful optical design and anti-reflection coated optics, our experiment could probably have reached $5\ 10^{-8}$ raw contrast. More importantly, we have demonstrated that with the coronagraph + wavefront control architecture adopted in our experiment, static speckles can be pushed down very low ($3\ 10^{-9}$) in long exposures. Our system successfully removed long term correlations in the coherent speckles, and their averaged level in long exposure was reduced with a $1/\sqrt{T}$ law. The combination of a high performance PIAA coronagraph and a focal-plane based wavefront control therefore appears extremely attractive for ground-based Extreme-AO. In that regard, our experiment has been a successful validation of the key technologies and control algorithms of the Subaru Corona-

graphic Extreme-AO (SCExAO) system currently in assembly. The major differences between the SCExAO PIAA coronagraph and our laboratory prototype are (1) the need to design and operate a PIAA coronagraph on a centrally obscured pupil with thick spider vanes and (2) the need for corrective optics to recover a wide field of view. These two requirements have been validated in a separate laboratory experiment using the final SCExAO coronagraph optics (Lozi et al. 2009).

Our experiment was limited at the $2\ 10^{-7}$ contrast by an optical ghost and at the $5\ 10^{-8}$ contrast by turbulence or vibrations. The PIAA coronagraph could therefore not be tested to the contrast level required for direct imaging of Earthlike planets from space (approximately 10^{-10}), although several key concepts were demonstrated, including simultaneous operation of a low-order wavefront sensor using starlight in the PSF core and high-order wavefront sensor using scattered light in the science focal plane. New calibration schemes which will be very useful for high contrast coronagraphy were also developed and validated, such as the system response matrix optimization loop, which can slowly run in the "background" to fine-tune the system.

PIAA coronagraph technologies for high contrast space applications are now being actively developed and tested at NASA Ames Research Center and NASA Jet Propulsion Laboratory. A new set of PIAA mirrors was recently manufactured to higher surface accuracy than the ones we used, and is being integrated within the High Contrast Imaging Testbed (HCIT) vacuum chamber at NASA Jet Propulsion Laboratory. We note that the HCIT chamber has already demonstrated stability to the 1e-10 contrast with a Lyot-type coronagraph (Trauger & Traub 2007). The exper-

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iment described in this paper served as a precursor to this new step, which is aimed at reaching higher contrast (minimum goal of 1e-9) in broadband light using a two-DM wavefront correction. In parallel to this effort, a highly flexible high stability air testbed at NASA Ames Research Center is coming online to explore technology and architecture trades for PIAA systems.

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